CORRIGENDA

to

Spectral Atlas for Amateur Astronomers...

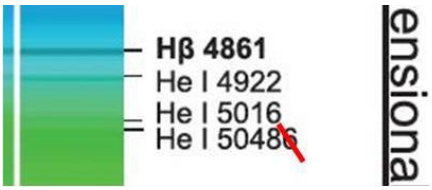
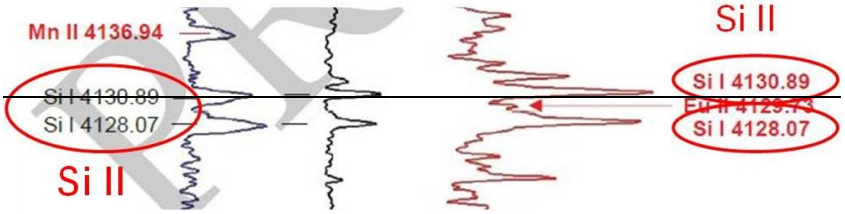
and

Spectroscopy for Amateur Astronomers...

Version 3.0

18 November 2025

Spectral Atlas for Amateur Astronomers

Page No.	Chapter/ Plate	Correction
21	Plate 1	
121	Plate 44	
74	14.5	lines of the H-Balmer series show up here, however, Hα is significantly dampened. In the Mira profile only Hγ and Hδ are in emission. At this low resolution the identification of the
144	25.7	all conditions are met, the degenerate electron gas of the iron core can no longer withstand the gravitational pressure. In contrast to the core-collapse scenario the stellar core consists here mostly of reactive carbon and oxygen, which is why the object immediately explodes. Therefore the SN type Ia is
87	16.7	broadened hydrogen lines no further spectral features are recognizable here. The FWHM equivalent width of the Hβ line is FWHM EW ≈ 65 Å, about 2.4 times as large as the impressive H-absorptions of the main sequence star Sirius, A1Vm (FWHM EW ≈ 27 Å). The classification digit 2.3 suggests here an effective temperature of about 22,000 K. Recording information: DADOS 50 µm slit, C8, Atik 314L+: 1×1800 s, 2×2 binning. WD 0413–077 40 Eridani B This white dwarf has spectral class DA 2.9, with m_v = 9.5 and is at a distance of 16 ly [11]. Here again, the strongly broadened hydrogen lines are the most prominent spectral features. The equivalent width of the Hβ line is here FWHM EW ≈ 78 Å even about three times as large as the main sequence star

<table>
<tr><td>178</td><td>28.12</td><td>5006.8, N I 5198.5, He II 5411.5
~~[Cl III]~~ [CIII] 5517.2, ~~[Cl III]~~ [CIII] 5537.7, [O I] 5577.4, [N II] 5754.8, He I 5875.6, [O I] 6300.2, [S III] 6310.2</td></tr>
<tr><td>144</td><td>Table 25.1</td><td>Table 25.1 Formation of heavy elements in stars with $M_i \approx 15\ M_\odot$
<table><tr><th>Stage</th><th>Fuel or Product</th><th>Ash or Product</th><th>Temperature [K]</th><th>Density [g cm^{-3}]</th><th>Time Scale [Years]</th></tr><tr><td>Hydrogen</td><td>H</td><td>He</td><td>3.5×10^7</td><td>5.8</td><td>$11 \times 10^{\text{X}6}$</td></tr><tr><td>Helium</td><td>He</td><td>C, O</td><td>1.8×10^8</td><td>1390</td><td>$2 \times 10^{\text{X}6}$</td></tr></table></td></tr>
<tr><td>193</td><td>Plate 72</td><td>Instead of the Blue Snowball Nebula NGC 7662, mistakenly the Saturn Nebula NGC 7009 is depicted.</td></tr>
</table>

Spectroscopy for Amateur Astronomers

Page No.	Chapter/ Plate	Correction
48	Table 5.4	150–2000 Wide field spectroscopic surveys e.g. by objective prisms General determination of spectral features e.g. emissions or absorptions General Stallar Spectral classification Spectral energy distribution (SED curves) Redshift of very faint quasars and galaxies ~~General stellar spectral classification~~ classification of faint novae and supernovae Excitation class of emission nebulae
68	Table 8.1	Methods of Calibration and Normalization Task: λ-calibration by rest wavelengths of known lines Relative measurement of a wavelength difference $\Delta\lambda$: ~~R~~ P
121	14.1.4	Inverted: Flare star Gliese 388, ~~LBV~~ AGB star Mira, S-type star R Cygni Irregular: Nova Delphini, recurrent nova T Crb, dwarf nova SS Cygni.